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LETTER TO THE EDITOR

## On the triple peaks of SNHunt248 in NGC 5806<sup>★</sup>

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### ABSTRACT

We present our findings on a supernova (SN) impostor, SNHunt248, based on optical and near-IR data spanning  $\sim 15$  yr before discovery, to  $\sim 1$  yr post-discovery. The light curve displays three distinct peaks, the brightest of which is at  $M_R \sim -15.0$  mag. The post-discovery evolution is consistent with the ejecta from the outburst interacting with two distinct regions of circumstellar material. The  $0.5\text{--}2.2\ \mu\text{m}$  spectral energy distribution at  $-740$  d is well-matched by a single 6700 K blackbody with  $\log(L/L_\odot) \sim 6.1$ . This temperature and luminosity support previous suggestions of a yellow hypergiant progenitor; however, we find it to be brighter than the brightest and most massive Galactic late-F to early-G spectral type hypergiants. Overall the historical light curve displays variability of up to  $\pm 1$  mag. At current epochs ( $\sim 1$  yr post-outburst), the absolute magnitude ( $M_R \sim -9$  mag) is just below the faintest observed historical absolute magnitude  $\sim 10$  yr before discovery.

**Key words.** stars: massive – stars: mass-loss – supernovae: general – supernovae: individual: SNHunt248

## 1. Introduction

Extragalactic transient events that exhibit fainter peak absolute magnitudes than most common types of SNe have been dubbed “SN impostors” (Van Dyk et al. 2000). Some of them have been linked to luminous blue variable (LBV) stars. By definition, these are non-terminal outbursts, thought to arise due to discrete mass-loss episodes via currently ill-understood mechanisms.

Here we report observations of a recently discovered eruptive transient in NGC 5806 ( $D_L = 22.5$  Mpc,  $\mu = 31.76$  mag, Tully et al. 2009). Originally discovered by CRTS (Drake et al. 2009) on 2014 May 21.18 UT,  $a \gtrsim 1$  mag brightening was reported shortly thereafter (Zheng et al. 2014), together with a spectroscopic similarity to other SN impostors. Early observations, reported by Mauerhan et al. (2015), show two light curve phases (“2014a” and “2014b”) with a slow rise in the 2014a phase followed by a rapid brightening to the 2014b peak. We adopt the beginning of this latter phase at JD = 2 456 812.35 (2014 June 3.85 UT) as our reference epoch. Mauerhan et al. (2015) identified a precursor at the location of SNHunt248 ( $\alpha = 14^{\text{h}}59^{\text{m}}59^{\text{s}}.47$  and  $\delta = +01^{\circ}54'26''.6$ , J2000) in archival HST data

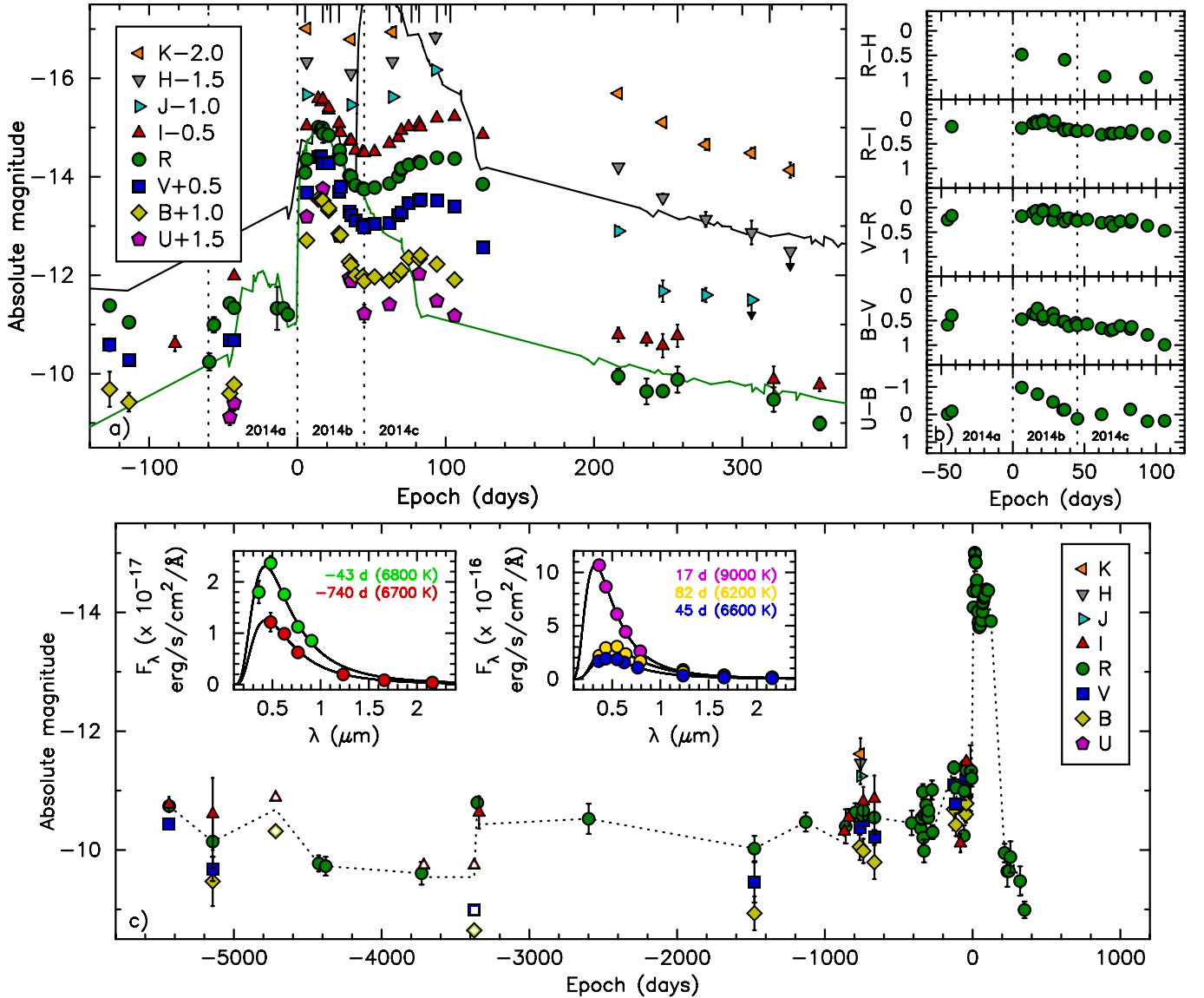
that led them to conclude that the star was of the cool hypergiant variety.

In what follows, we present archival data of SNHunt248 over a 15 yr period before discovery and our post-discovery photometric and spectroscopic follow-up out to about 1 yr.

## 2. Observations

Following the reported discovery of SNHunt248, we began a follow-up campaign using several optical and near-IR facilities (see Tables 1 to 4). We recovered available historical light curve information on the precursor of SNHunt248 from publicly-available archival images of the field. Standard IRAF-based procedures were followed for bias subtraction and flat-field correction of the optical follow-up images and carried out using primarily the QUBA pipeline (Valenti et al. 2011). Processing of the near-IR follow-up frames included additional steps involving sky subtraction, image alignment and stacking using a slightly modified version of the IRAF based NOTCAM package. PSF fitting photometry was carried out within the QUBA pipeline, which also interpolates the sky background at the transient location based on the surrounding environment. The optical photometry was calibrated based on the SDSS DR9 *ugriz* magnitudes of 15 field stars. For the Johnson-Cousins filters,

<sup>★</sup> Tables 1–4 are available in electronic form at <http://www.aanda.org>



**Fig. 1.** a) Absolute light curves of SNHunt248. The solid vertical lines indicate the epochs of spectroscopic observations. For comparison, we show the absolute  $R$ -band light curve of SN 2009ip (Pastorello et al. 2013; Fraser et al. 2013, 2015, black curve: JD = 2456155 corresponds to day 0, no shift in magnitude; and green curve: day 0 epoch JD = 2456195, +3.0 mag shift) emphasizing the similar evolution between some of the key phases. b) Colour evolution of SNHunt248. c) Historical light curve of SNHunt248. From day 0 onwards only the  $R$ -band points are included for clarity. The HST archival data points from Mauerhan et al. (2015) are shown with open symbols. The  $R$ -band points are connected with a dotted line to guide the eye (for epochs with an  $I$ -band point but a missing  $R$ -band magnitude, the latter is very roughly estimated assuming  $R - I = 0.2$  mag, corresponding to the  $-740$  d spectral energy distribution (SED)). Insets display blackbody fits to selected pre- and post-discovery epochs of observed data (interpolated post-discovery near-IR data are included for comparison).

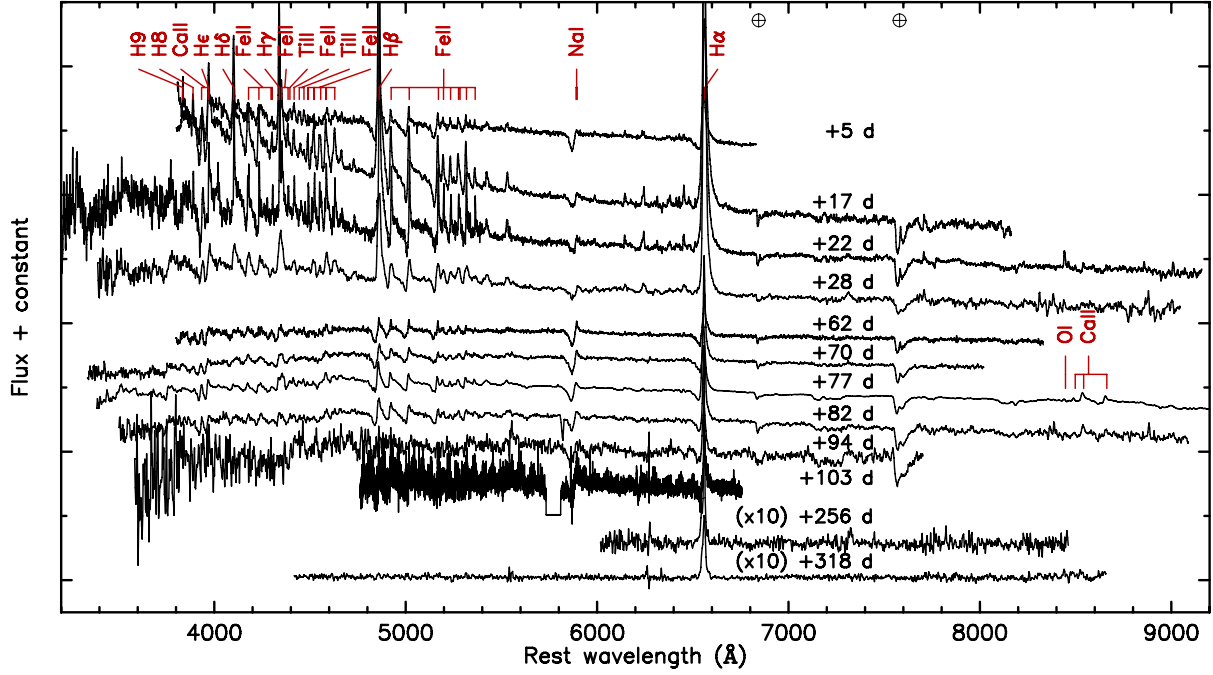
we used the transformations of Jester et al. (2005) to convert the sequence star magnitudes to the  $UBVRI$  system. The near-IR images were calibrated using 2MASS  $JHK$  magnitudes of 10 sequence stars. Average instrument-specific colour terms were derived from standard star observations and applied only for optical follow up, otherwise zero colour terms were assumed. The resulting photometry is presented in Tables 1 to 3 and in Fig. 1.

Spectroscopic observations from the NOT, WHT, and TNG were also reduced using the QUBA pipeline and included the standard steps of bias subtraction, flat field correction, spectrum extraction, wavelength and relative flux calibration. Wavelength calibration was derived based on arc lamp exposures. Relative flux calibration was carried out using a sensitivity curve derived with observations of spectrophotometric standard stars obtained on the same night and setup as SNHunt248. The GTC+OSIRIS

and VLT+FORS2 spectra were reduced in a similar way using basic IRAF tasks. The VLT+UVES spectrum was reduced using the Reflex-based pipeline workflow. Furthermore, all the spectra were absolute flux calibrated based on the optical photometry of SNHunt248. The spectroscopic observations are summarized in Table 4 and the spectral sequence of SNHunt248 is shown in Fig. 2.

### 3. Analysis

We adopt  $A_V^{\text{Gal}} = 0.140$  mag for Galactic extinction towards SNHunt248 (Schlafly & Finkbeiner 2011). In our highest resolution spectrum (+103 d,  $R = 26000$ ), we do not detect the narrow interstellar medium (ISM) absorption components of the Na I D at the expected redshift of NGC 5806. SNHunt248 lies at



**Fig. 2.** Spectral time series of SNHunt248. The spectra have been dereddened, redshift corrected to the rest frame and vertically shifted for clarity. The strongest telluric features are indicated by  $\oplus$  symbols. The identification of the most prominent lines is based on that of SN 2009kn (see Kankare et al. 2012, and references therein).

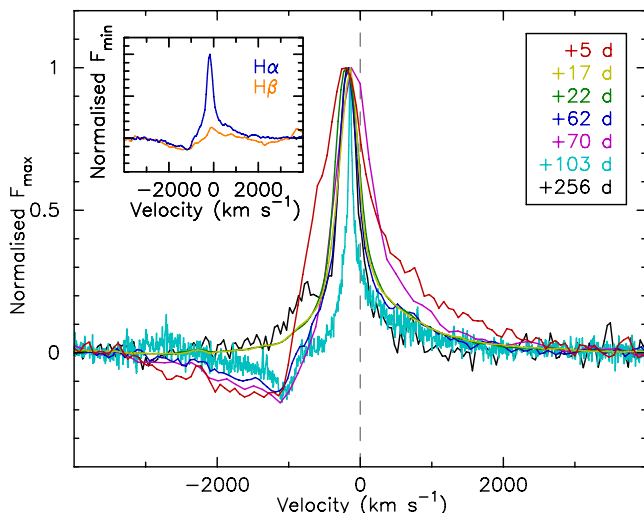
a relatively large projected offset ( $\sim 6.4$  kpc) from the host nucleus in an uncrowded region. This is suggestive of small host galaxy reddening and we adopt  $A_V^{\text{host}} = 0$  mag.

The most striking feature in the post-discovery evolution of SNHunt248 is its triple-peaked light curve. Our data shows a clear decline in the  $R$ -band light curve between  $-127$  and  $-60$  d which marks the beginning of the rise to the phase 2014a, with the first peak at  $M_R \approx -11.4$  mag. In the 2014b phase the light curve rises  $>3$  mag to the second peak (and maximum) in  $\sim 15$  days in all the optical bands, followed by a decline of  $\sim 1$  mag for the next  $\sim 30$  days. This is followed by a phase of re-brightening. The rise to the third maximum is less rapid and takes  $\sim 50$  d, with the blue bands evolving faster than the red ones. Extending the naming convention of Mauerhan et al. (2015), we dub this third light curve phase as “2014c”. The absolute magnitude of this peak ( $M_R \approx -14.4$  mag) does not reach the brightness of the 2014b maximum ( $M_R \approx -15.0$  mag) which is somewhat brighter than the light curve maximum of most events classified as SN impostors (Smith et al. 2011), but fainter than normal SNe. SNHunt248 colour is bluest at the peak of the phase 2014b peak, followed by a fairly linear evolution towards a redder colour (see Fig. 1). By  $\sim +250$  d SNHunt248 has already reached the historical light curve minima observed  $\sim 3000$  to  $7000$  d before the outburst. Overall the historical light curve of SNHunt248 has shown variability of up to  $\pm 1$  mag over the last 1.5 decades.

Our earliest spectrum (+5 d, phase 2014b) shows P Cygni line profiles of Balmer lines and a forest of Fe II and other metal lines with a strong absorption. The Balmer lines are the most prominent features in the spectra of SNHunt248. The  $H\alpha$  absorption minimum indicates a velocity of  $\sim 1100$  km  $s^{-1}$  while the blue edge of the asymmetric absorption component extends to  $\sim 3000$  km  $s^{-1}$ . No strong evolution of the metal lines is seen whereas the most notable evolution in the line profiles is the disappearance of the absorption features in the Balmer lines by the time the light curve reaches the phase 2014b maximum. The

absorption feature in the Balmer lines remains absent throughout the light curve decline of this phase. However, as the light curve evolves into the rise towards phase 2014c, the absorption features reappear in the Balmer lines (Fig. 3). It is noteworthy that the absorption component at +62 and +70 d, i.e. in phase 2014c, displays an essentially identical absorption component profile compared to the early +5 d spectrum, indicating material moving at the same velocity as in phase 2014b. Unlike the Balmer lines, the metal lines show a P Cygni profile throughout our observed time series from the early 2014b phase to the 2014c peak, but become less prominent and in the late-time spectra the low signal-to-noise ratio prevents identification of all but the strongest lines.

The field of SNHunt248 was observed  $\sim 2$  yr before the 2014 discovery with WHT+LIRIS in  $JHK$  and GTC+OSIRIS in  $gri$  within 19 days. The SED is well described by a single black-body yielding  $T_{\text{BB}} \approx 6700$  K,  $R_{\text{BB}} \approx 5.9 \times 10^{13}$  cm ( $\sim 850 R_{\odot}$ ), and  $L_{\text{BB}} \approx 5.0 \times 10^{39}$  erg  $s^{-1}$  ( $\sim 1.3 \times 10^6 L_{\odot}$ ), see Fig. 1. This would suggest the SNHunt248 progenitor to be a yellow hypergiant (YHG) close to the “Yellow Void” (de Jager 1998). YHGs are unstable and eruptive stars, and some of them have dusty circumstellar envelopes. The SNHunt248 SED shows no sign of a near-IR excess at day  $-740$  out to  $2.2 \mu\text{m}$ . Furthermore, the variable luminosity of SNHunt248 at  $-740$  d is above, albeit only briefly, the Humphreys-Davidson limit ( $\log L \approx 5.8 L_{\odot}$  at the corresponding  $T$ ) where massive LBVs in outburst are often found. In the coolest eruptive states, LBVs can resemble type F hypergiants with temperatures of  $\sim 7000$ – $8000$  K (Humphreys & Davidson 1994). The SNHunt248 precursor is close, but below this temperature lower limit. For comparison, Mauerhan et al. (2015) estimated  $T \approx 6500$  K,  $R \approx 500 R_{\odot}$ , and  $L \approx 4 \times 10^5 L_{\odot}$  for the SNHunt248 precursor on  $-3373$  d. This is roughly the observed range of the pre-discovery variability, prominent mainly in the radius and luminosity. It is possible that the observed variability in the historical light curve is arising some or all the



**Fig. 3.** Selection of continuum-subtracted and peak-normalised  $H\alpha$  profiles of SNHunt248. For clarity, only spectra with  $R > 500$  are shown. The inset shows the minimum-normalised +62 d  $H\alpha$  and  $H\beta$  profiles with identical absorption components. The narrow emission component is resolved only in the +103 d VLT+UVES spectrum.

times from discrete mass-loss episodes, possibly related to the formation of a pseudo-photosphere.

Based on our early photometry, the initial outburst of SNHunt248, and the rise to phase 2014a, commenced at  $\sim -60$  d, followed by a brightening of the light curves to the double-peaked maximum powered by strong circumstellar medium (CSM) interaction. This is consistent with the early ( $-9$  d) spectrum, reported by Mauerhan et al. (2015), displaying similar P Cygni velocities *before* phase 2014b as seen during phase 2014b and 2014c, since the P Cygni absorption component likely arise from the outburst. It is therefore also unlikely that the 2014c phase arose from a new major outburst of the progenitor since no evolution is observed in the velocities of the  $H\alpha$  absorption component. A more likely explanation is that the outburst ejecta from the 2014a event collided into two shell-like regions of dense CSM, powering phases 2014b and 2014c. A possible Galactic YHG analogue to this scenario would be IRAS 17163-3907 which has been shown to have two shell-like structures at  $3.6 \times 10^{16}$  and  $9 \times 10^{16}$  cm (Lagadec et al. 2011). Differences in the light curve and spectral evolution of phases 2014b and 2014c likely arise because the CSM shells are not identical in terms of their masses, densities, and velocities. The narrow emission component of  $H\alpha$  is resolved in our +103 d (phase 2014c) spectrum; a Lorentzian fit yields a FWHM of  $\sim 90$  km  $s^{-1}$ , which is consistent with stellar winds of many LBVs and also some YHGs (Lagadec et al. 2011).

We assume an initial radius of  $6 \times 10^{13}$  cm from our  $-740$  d data, the initial eruption already taking place at  $-60$  d with a bulk expansion velocity of  $\sim 1100$  km  $s^{-1}$ . The rise to the next maximum (phase 2014b) results in  $\sim 6.4$  to  $\sim 7.7 \times 10^{14}$  cm and phase 2014c maximum from  $\sim 1.1$  to  $\sim 1.5 \times 10^{15}$  cm. During the 2014b phase we do not resolve the narrow emission component; however, the resolution of our WHT spectrum allows us to set an upper limit of  $< 300$  km  $s^{-1}$  for the CSM wind, suggesting a period of mass-loss  $\geq 0.7$  yr before the outburst. For phase 2014c, the  $90$  km  $s^{-1}$  CSM wind suggest the earlier mass-loss episode to have taken place from  $\sim 3.5$  to  $\sim 5$  yr before the events of 2014.

## 4. Conclusions

We find the unique 2014 triple-peaked brightening of SNHunt248 to be arising from an outburst of a progenitor with an appearance of a very luminous YHG with no clear near-IR excess. With an optical peak magnitude of roughly  $-15$  mag, the SNHunt248 event is among the most luminous SN impostors. The historical light curve of SNHunt248 presented here offers a rare pre-discovery view to the evolution of an SN impostor precursor and shows variability episodes. The photometric and spectroscopic evolution of SNHunt248 is consistent with the outburst ejecta interacting with two distinct regions of CSM giving rise to a luminous display, not observed in Galactic YHGs.

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**Table 1.** Optical *UBVRI* photometry for SNHunt248 with the errors given in brackets.

JD (2 400 000+)	<i>t</i> (d)	<i>m<sub>U</sub></i> (mag)	<i>m<sub>B</sub></i> (mag)	<i>m<sub>V</sub></i> (mag)	<i>m<sub>R</sub></i> (mag)	<i>m<sub>I</sub></i> (mag)	Telescope+instrument	Seeing ( $''$ )
51 373.79	-5439	...	...	21.458(0.053)	21.138(0.058)	21.076(0.130)	CFHT+CFH12K	0.9
51 668.89	-5143	...	22.475(0.416)	22.219(0.204)	21.736(0.442)	21.237(0.608)	SDSS <sup>a</sup>	1.0
52 386.06	-4426	...	...	...	22.104(0.135)	...	CFHT+CFH12K	1.2
52 433.98	-4378	...	...	...	22.150(0.157)	...	CFHT+CFH12K	1.2
53 082.70	-3730	...	...	...	22.270(0.190)	...	INT+WFC <sup>b</sup>	1.2
53 459.66	-3353	...	...	...	21.082(0.103)	...	INT+WFC	1.2
53 471.81	-3341	...	...	...	...	21.212(0.272)	CTIO4m+Mosaic-2	1.6
54 213.81	-2599	...	...	...	21.352(0.254)	...	KPNO4m+Mosaic-1	1.9
55 335.77	-1477	...	23.013(0.285)	22.445(0.339)	21.854(0.214)	...	KPNO4m+Mosaic-1	2.8
55 682.95	-1129	...	...	...	21.407(0.159)	...	PS1+GPC1 <sup>c</sup>	1.3
55 944.65	-868	...	...	...	...	21.528(0.077)	Gemini-N+GMOS <sup>d</sup>	0.7
55 953.70	-859	...	...	...	21.481(0.288)	...	LT+RATCam <sup>b</sup>	1.2
55 973.86	-838	...	...	...	...	21.296(0.144)	NTT+EFOSC2	0.9
56 019.92	-792	...	...	...	21.246(0.154)	...	FTN+FS02	2.1
56 048.57	-764	...	21.893(0.223)	21.523(0.121)	...	...	NOT+ALFOSC <sup>a</sup>	0.9
56 065.89	-746	...	...	...	21.291(0.083)	...	PS1+GPC1 <sup>c</sup>	0.9
56 072.64	-740	...	21.963(0.205)	21.405(0.089)	21.215(0.179)	21.025(0.244)	GTC+OSIRIS <sup>a</sup>	1.0
56 147.54	-665	...	22.153(0.283)	21.675(0.112)	21.332(0.280)	20.980(0.390)	NTT+EFOSC2 <sup>a</sup>	1.3
56 400.03	-412	...	...	...	21.421(0.207)	...	PS1+GPC1 <sup>c</sup>	1.5
56 461.44	-351	...	...	...	21.514(0.176)	...	TNG+DOLORES	1.1
56 466.50	-346	...	...	...	21.339(0.117)	...	VLT+FOR2 <sup>e</sup>	0.8
56 475.89	-336	...	...	...	20.899(0.130)	...	FTN+FS02	2.4
56 477.49	-335	...	...	...	21.671(0.087)	...	VLT+FOR2 <sup>e</sup>	0.5
56 483.57	-329	...	...	...	21.893(0.192)	...	VLT+FOR2 <sup>e</sup>	0.9
56 485.56	-327	...	...	...	21.287(0.063)	...	VLT+FOR2 <sup>e</sup>	0.5
56 497.47	-315	...	...	...	21.120(0.064)	...	VLT+FOR2 <sup>e</sup>	0.8
56 501.50	-311	...	...	...	21.287(0.135)	...	VLT+FOR2 <sup>e</sup>	1.0
56 512.52	-300	...	...	...	21.331(0.104)	...	VLT+FOR2 <sup>e</sup>	0.9
56 513.49	-299	...	...	...	21.217(0.081)	...	VLT+FOR2 <sup>e</sup>	0.7
56 537.48	-275	...	...	...	20.868(0.162)	...	VLT+FOR2 <sup>e</sup>	1.4
56 541.47	-271	...	...	...	21.576(0.117)	...	VLT+FOR2 <sup>e</sup>	0.7
56 685.73	-127	...	21.261(0.356)	20.809(0.128)	20.490(0.104)	...	NOT+ALFOSC <sup>a,b</sup>	1.4
56 698.72	-114	...	21.523(0.190)	21.121(0.076)	20.830(0.083)	...	NOT+ALFOSC <sup>a,b</sup>	1.0
56 729.81	-83	...	...	...	...	21.728(0.155)	VST+OmegaCAM <sup>d</sup>	0.6
56 752.95	-59	...	...	...	21.636(0.173)	...	PS1+GPC1 <sup>c</sup>	1.2
56 756.01	-56	...	...	...	20.883(0.156)	...	PS1+GPC1 <sup>c</sup>	1.5
56 766.79	-46	21.362(0.161)	21.348(0.091)	20.715(0.045)	20.447(0.047)	...	SOAR+Goodman	1.2
56 769.68	-43	21.081(0.150)	21.164(0.086)	20.722(0.032)	20.538(0.053)	20.356(0.069)	CTIO4m+DECam <sup>a</sup>	1.0
56 798.68	-14	...	...	...	20.550(0.436)	...	MLS <sup>f</sup>	2.5
56 802.91	-9	...	...	...	20.549(0.104)	...	PS1+GPC1 <sup>c</sup>	1.4
56 805.85	-7	...	...	...	20.672(0.137)	...	PS1+GPC1 <sup>c</sup>	1.5
56 817.47	5	...	...	...	17.794(0.021)	...	NOT+ALFOSC	0.6
56 818.60	6	17.292(0.107)	18.237(0.071)	17.718(0.038)	17.523(0.032)	17.310(0.025)	NOT+StanCam	0.9
56 826.39	14	...	17.399(0.027)	16.994(0.011)	16.871(0.025)	16.754(0.034)	LT+IO:O <sup>a</sup>	0.9
56 828.40	16	...	17.410(0.023)	16.988(0.010)	16.900(0.018)	16.820(0.023)	LT+IO:O <sup>a</sup>	0.9
56 829.49	17	16.724(0.030)	17.428(0.012)	17.129(0.012)	16.887(0.011)	16.770(0.020)	NOT+ALFOSC	1.2
56 829.68	17	...	...	...	17.002(0.184)	...	MLS <sup>f</sup>	2.3
56 833.40	21	...	17.628(0.037)	17.106(0.013)	17.043(0.019)	16.989(0.024)	LT+IO:O <sup>a</sup>	1.4
56 833.41	21	...	17.582(0.016)	17.126(0.007)	17.031(0.014)	16.943(0.019)	WHT+ACAM <sup>a</sup>	1.5
56 840.44	28	17.680(0.032)	18.101(0.027)	17.700(0.017)	17.422(0.014)	17.263(0.020)	NOT+ALFOSC	1.2
56 840.80	28	...	...	...	17.337(0.016)	...	PS1+GPC1 <sup>c</sup>	1.9

**Notes.** Seeing is derived from the image most similar to *R*-band, and corresponds to the average FWHM of a selection of sequence stars. CFHT = 3.58-m Canada-France-Hawaii Telescope (Mauna Kea, Hawaii); SDSS = 2.5-m Sloan Digital Sky Survey telescope (Apache Point Observatory, New Mexico); INT = 2.5-m *Isaac Newton* Telescope (La Palma, Spain); CTIO4m = 4-m Blanco Telescope (Cerro Tololo, Chile); KPNO4m = 4-m Mayall Telescope (Kitt Peak, Arizona); PS1 = 1.8-m Panoramic Survey Telescope & Rapid Response System 1 (Haleakala, Hawaii); Gemini-N = 8.1-m Gemini North Telescope (Mauna Kea, Hawaii); LT = 2.0-m Liverpool Telescope (La Palma, Spain); NTT = 3.58-m ESO New Technology Telescope (La Silla, Chile); FTN = 2.0-m Faulkes Telescope North (Haleakala, Hawaii); NOT = 2.56-m Nordic Optical Telescope (La Palma, Spain); GTC = 10.4-m Gran Telescopio Canarias (La Palma, Spain); TNG = 3.58-m Telescopio Nazionale *Galileo* (La Palma, Spain); VLT = 8.2-m Very Large Telescope (Cerro Paranal, Chile); VST = 2.6-m VLT Survey Telescope (Cerro Paranal, Chile); SOAR = 4.1-m Southern Astrophysical Research Telescope (Cerro Pachón, Chile); MLS = 1.5-m Mt. Lemmon Survey Telescope (Steward Observatory, Arizona); AS1.82-m = 1.82-m Copernico Telescope (Asiago, Italy); WHT = 4.2-m *William Herschel* Telescope (La Palma, Spain). <sup>(a)</sup> Table 2 *ugri* magnitudes converted to the *UBVRI* system using the transformations of Jester et al. (2005). <sup>(b)</sup> *r*-band image calibrated directly into *R*. <sup>(c)</sup> PS1 *w<sub>p1</sub>* image calibrated directly into *R*. <sup>(d)</sup> *i*-band image calibrated directly into *I*. <sup>(e)</sup> VLT blocking filter GG435 image calibrated directly into *R*. <sup>(f)</sup> Unfiltered image calibrated directly into *R*.

Table 1. continued.

JD (2 400 000+)	<i>t</i> (d)	<i>m<sub>U</sub></i> (mag)	<i>m<sub>B</sub></i> (mag)	<i>m<sub>V</sub></i> (mag)	<i>m<sub>R</sub></i> (mag)	<i>m<sub>I</sub></i> (mag)	Telescope+instrument	Seeing (")
56 841.41	29	...	18.127(0.030)	17.603(0.013)	17.519(0.025)	17.443(0.033)	LT+IO:O <sup>a</sup>	1.7
56 847.42	35	18.539(0.070)	18.671(0.062)	18.102(0.022)	17.850(0.024)	17.622(0.032)	NOT+ALFOSC	1.6
56 848.54	36	18.602(0.081)	18.741(0.039)	18.175(0.021)	17.868(0.020)	17.605(0.051)	NOT+StanCam	1.3
56 851.41	39	...	18.948(0.045)	18.288(0.020)	18.049(0.039)	17.807(0.053)	LT+IO:O <sup>a</sup>	0.6
56 856.39	44	...	18.974(0.027)	18.378(0.011)	18.110(0.027)	17.839(0.038)	LT+IO:O <sup>a</sup>	0.9
56 857.42	45	19.258(0.181)	19.068(0.040)	18.423(0.023)	18.131(0.037)	17.855(0.046)	AS1.82m+AFOSC <sup>a</sup>	1.8
56 864.39	52	...	18.979(0.043)	18.359(0.017)	18.098(0.032)	17.834(0.043)	LT+IO:O <sup>a</sup>	0.7
56 874.37	62	19.072(0.049)	19.048(0.026)	18.346(0.015)	18.015(0.014)	17.668(0.020)	NOT+ALFOSC	1.1
56 880.41	68	...	18.933(0.082)	18.185(0.031)	17.870(0.047)	17.547(0.060)	LT+IO:O <sup>a</sup>	1.2
56 882.43	70	...	18.850(0.031)	18.123(0.013)	17.728(0.013)	17.402(0.015)	TNG+DOLORES	1.5
56 882.50	70	...	...	...	17.709(0.012)	...	CTIO4m+DECam <sup>b</sup>	1.0
56 887.38	75	...	18.592(0.030)	17.937(0.013)	17.636(0.023)	17.329(0.031)	LT+IO:O <sup>a</sup>	1.0
56 894.41	82	18.452(0.045)	18.609(0.020)	17.888(0.014)	17.575(0.015)	17.255(0.023)	NOT+ALFOSC	1.0
56 895.37	83	...	18.539(0.030)	17.866(0.014)	17.600(0.026)	17.331(0.036)	LT+IO:O <sup>a</sup>	0.9
56 906.36	94	19.000(0.065)	18.721(0.020)	17.881(0.016)	17.490(0.012)	17.152(0.014)	NOT+ALFOSC	1.0
56 918.35	106	19.301(0.063)	19.040(0.025)	18.001(0.014)	17.509(0.015)	17.120(0.019)	NOT+ALFOSC	1.2
56 937.32	125	...	...	18.834(0.077)	18.026(0.033)	17.486(0.029)	NOT+StanCam	1.2
57 028.75	216	...	...	...	21.930(0.156)	21.550(0.146)	NOT+StanCam	0.6
57 047.78	235	...	...	...	22.236(0.261)	21.643(0.128)	NOT+ALFOSC	0.9
57 058.71	246	...	...	...	22.231(0.085)	21.773(0.239)	NOT+StanCam	0.6
57 068.65	256	...	...	...	21.996(0.264)	21.572(0.227)	WHT+ACAM	1.3
57 133.49	321	...	...	...	22.400(0.247)	22.461(0.267)	WHT+ACAM	1.1
57 164.53	352	...	...	...	22.885(0.139)	22.568(0.129)	VLT+FORIS2	0.9

Table 2. Original *ugriz* photometry for SNHunt248 with the errors given in brackets.

JD (24 00000+)	<i>t</i> (d)	<i>m<sub>u</sub></i> (mag)	<i>m<sub>g</sub></i> (mag)	<i>m<sub>r</sub></i> (mag)	<i>m<sub>i</sub></i> (mag)	<i>m<sub>z</sub></i> (mag)	Telescope+instrument	Seeing (")
51 668.89	-5143	...	22.258(0.305)	22.208(0.274)	21.924(0.303)	...	SDSS	1.0
56 048.57	-764	...	21.635(0.162)	21.459(0.173)	...	...	NOT+ALFOSC	0.9
56 072.64	-740	...	21.637(0.152)	21.255(0.106)	21.275(0.123)	...	GTC+OSIRIS	1.0
56 147.54	-665	...	21.856(0.211)	21.562(0.117)	21.423(0.240)	...	NTT+EFOSC2	1.3
56 385.88	-426	...	...	...	...	21.402(0.050)	CTIO4m+DECam	1.0
56 685.73	-127	...	20.973(0.266)	20.707(0.106)	...	...	NOT+ALFOSC	1.4
56 698.72	-114	...	21.253(0.141)	21.042(0.083)	...	...	NOT+ALFOSC	1.0
56 769.68	-43	21.915(0.146)	20.880(0.064)	20.625(0.030)	20.652(0.032)	20.533(0.072)	CTIO4m+DECam	1.0
56 826.39	14	...	17.128(0.020)	16.914(0.013)	17.005(0.019)	...	LT+IO:O	0.9
56 828.40	16	...	17.133(0.017)	16.901(0.012)	17.028(0.009)	...	LT+IO:O	0.9
56 833.40	21	...	17.315(0.028)	16.972(0.011)	17.125(0.009)	...	LT+IO:O	1.4
56 833.41	21	...	17.293(0.012)	17.022(0.008)	17.142(0.010)	...	WHT+ACAM	1.5
56 841.41	29	...	17.813(0.022)	17.469(0.016)	17.600(0.015)	...	LT+IO:O	1.7
56 851.41	39	...	18.585(0.033)	18.091(0.024)	18.060(0.026)	...	LT+IO:O	0.6
56 856.39	44	...	18.634(0.020)	18.210(0.013)	18.150(0.022)	...	LT+IO:O	0.9
56 857.42	45	19.522(0.195)	...	18.269(0.019)	18.203(0.020)	...	AS1.82m+AFOSC	1.8
56 864.39	52	...	18.630(0.032)	18.180(0.019)	18.127(0.021)	...	LT+IO:O	0.7
56 880.41	68	...	18.538(0.061)	17.947(0.029)	17.837(0.023)	...	LT+IO:O	1.2
56 881.50	69	...	...	...	...	17.767(0.021)	CTIO4m+DECam	1.2
56 887.38	75	...	18.231(0.022)	17.742(0.015)	17.647(0.013)	...	LT+IO:O	1.0
56 895.37	83	...	18.171(0.022)	17.662(0.018)	17.604(0.015)	...	LT+IO:O	0.9

Notes. Seeing is derived from the image most similar to *r*-band. Telescopes are detailed in Table 1.

**Table 3.** Near-IR *JHK* photometry for SNHunt248 with the errors given in brackets.

JD (2 400 000+)	<i>t</i> (d)	<i>m<sub>J</sub></i> (mag)	<i>m<sub>H</sub></i> (mag)	<i>m<sub>K</sub></i> (mag)	Telescope+instrument	Seeing ( $''$ )
56 053.65	-759	20.555(0.135)	20.321(0.237)	20.159(0.264)	WHT+LIRIS	1.0
56 818.58	6	17.129(0.039)	16.950(0.036)	16.764(0.057)	NOT+NOTCam	0.8
56 848.52	36	17.338(0.040)	17.189(0.041)	16.986(0.053)	NOT+NOTCam	0.9
56 876.44	64	17.177(0.048)	16.944(0.047)	16.835(0.095)	NOT+NOTCam	0.8
56 905.42	93	16.630(0.044)	16.451(0.091)	...	NOT+NOTCam	1.2
57 028.77	216	19.900(0.088)	19.088(0.056)	18.083(0.055)	NOT+NOTCam	0.7
57 058.76	246	21.121(0.219)	19.711(0.086)	18.669(0.073)	NOT+NOTCam	0.6
57 087.70	275	21.200(0.143)	20.151(0.151)	19.121(0.110)	NOT+NOTCam	0.7
57 118.58	306	>21.3	20.418(0.241)	19.292(0.100)	NOT+NOTCam	0.9
57 144.66	332	...	>20.8	19.640(0.153)	NOT+NOTCam	0.6

**Notes.** Seeing is derived from the image most similar to *K*-band. Telescopes are detailed in Table 1.

**Table 4.** Log of spectroscopic observations for SNHunt248.

JD (2 400 000+)	<i>t</i> (d)	Grism/grating	Slit ( $''$ )	<i>R</i> ( $\lambda/\Delta\lambda$ )	<i>R</i> ( $\text{km s}^{-1}$ )	Exp. time (s)	Telescope+instrument
56 817.5	5	#7	1.0	650	460	1800	NOT+ALFOSC
56 829.5	17	#7, #8	1.3	500, 770	600, 390	1800	NOT+ALFOSC
56 834.6	22	R300B, R158R	1.0	1100, 970	270, 310	900	WHT+ISIS
56 840.4	28	#4	1.3	270	1100	2400	NOT+ALFOSC
56 874.4	62	#7, #8	1.0	650, 1000	460, 300	2400	NOT+ALFOSC
56 882.4	70	LR-B	1.0	590	510	2700	TNG+DOLORES
56 889.3	77	R500B, R500R	1.0	320, 350	940, 860	1800	GTC+OSIRIS
56 894.4	82	#4	1.0	360	830	2400	NOT+ALFOSC
56 906.4	94	#4	1.3	270	1100	1800	NOT+ALFOSC
56 915.5	103	Red	1.5	26 000	12	2300	VLT+UVES
57 068.7	256	R316R	1.5	1300	230	5400	WHT+ISIS
57 130.8	318	300V	1.0	440	680	1800	VLT+FORS2

**Notes.** Telescopes are detailed in Table 1.