



THE UNIVERSITY *of* EDINBURGH

Edinburgh Research Explorer

Individual dynamical masses of DENIS J063001.4–184014AB reveal a likely young brown dwarf triple

Citation for published version:

Sahlmann, J, Dupuy, TJ, Burgasser, AJ, Filippazzo, JC, Martín, EL, Gagliuffi, DCB, Hsu, C, Lazorenko, PF & Liu, MC 2021, 'Individual dynamical masses of DENIS J063001.4–184014AB reveal a likely young brown dwarf triple', *Monthly Notices of the Royal Astronomical Society*, vol. 500, no. 4, pp. 5453-5461.
<https://doi.org/10.1093/mnras/staa3577>

Digital Object Identifier (DOI):

[10.1093/mnras/staa3577](https://doi.org/10.1093/mnras/staa3577)

Link:

[Link to publication record in Edinburgh Research Explorer](#)

Document Version:

Peer reviewed version

Published In:

Monthly Notices of the Royal Astronomical Society

General rights

Copyright for the publications made accessible via the Edinburgh Research Explorer is retained by the author(s) and / or other copyright owners and it is a condition of accessing these publications that users recognise and abide by the legal requirements associated with these rights.

Take down policy

The University of Edinburgh has made every reasonable effort to ensure that Edinburgh Research Explorer content complies with UK legislation. If you believe that the public display of this file breaches copyright please contact openaccess@ed.ac.uk providing details, and we will remove access to the work immediately and investigate your claim.



Individual Dynamical Masses of DENIS J063001.4–184014AB Reveal A Likely Young Brown Dwarf Triple^{*}

J. Sahlmann,^{1,2,3†} T. J. Dupuy,⁴ A. J. Burgasser,⁵ J. C. Filippazzo,³
 E. L. Martín,^{6,7,8} D. C. Bardalez Gagliuffi,⁹ C. Hsu,⁵ P. F. Lazorenko,¹⁰
 Michael C. Liu¹¹

¹*RHEA Group for the European Space Agency (ESA), European Space Astronomy Centre (ESAC), Camino Bajo del Castillo s/n, 28692 Villanueva de la Cañada, Madrid, Spain*

²*Independent researcher*

³*Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA*

⁴*Institute for Astronomy, University of Edinburgh, Royal Observatory, Blackford Hill, Edinburgh, EH9 3HJ, UK*

⁵*Center for Astrophysics and Space Science, University of California San Diego, La Jolla, CA, 92093, USA*

⁶*Instituto de Astrofísica de Canarias (IAC), Calle Vía Láctea s/n, E-38200 La Laguna, Tenerife, Spain*

⁷*Departamento de Astrofísica, Universidad de La Laguna (ULL), E-38206 La Laguna, Tenerife, Spain*

⁸*Consejo Superior de Investigaciones Científicas, E-28006 Madrid, Spain*

⁹*Department of Astrophysics, American Museum of Natural History, Central Park West at 79th Street, NY 10024, USA*

¹⁰*Main Astronomical Observatory, National Academy of Sciences of the Ukraine, Zabolotnogo 27, 03680 Kyiv, Ukraine*

¹¹*Institute for Astronomy, University of Hawaii, 2680 Woodlawn Drive, Honolulu, HI 96822, USA*

Accepted 2020 November 9. Received 2020 October 9; in original form 2020 August 5.

ABSTRACT

The binary nature of the M8.5 dwarf DENIS J063001.4–184014AB (DE0630–18) was discovered with astrometric monitoring from the ground, which determined the unresolved photocentric orbit and the trigonometric parallax of the system. Here we present radial-velocity monitoring and resolved observations in the near-infrared with Keck aperture masking that allow us to measure the system’s relative separation and brightness. By combining all available information, we determine the individual dynamical masses of the binary components to be $M_1 = 0.052_{-0.008}^{+0.009} M_{\text{Sun}}$ and $M_2 = 0.052_{-0.004}^{+0.005} M_{\text{Sun}}$, both firmly in the substellar regime. These masses are surprising given the object’s M8.5 optical spectral type and equivalent absolute magnitude, and the significant difference in brightness between the components ($\Delta K = 1.74 \pm 0.06$ mag). Our results suggest that DE0630–18 is a relatively young system (~ 200 Myr) with a secondary component that is itself a potentially unresolved binary.

Key words: brown dwarfs – astrometry – binaries: visual – parallaxes – stars: individual: DE0630–18

1 INTRODUCTION

The masses of ultracool dwarfs, i.e. very low-mass stars and brown dwarfs, are difficult to measure precisely; yet they are necessary to refine our theoretical understanding of these objects and their physics (e.g. Zapatero Osorio et al. 2004; Liu et al. 2008). The monitoring of binary star motions that are governed by gravitational interaction give us the opportunity to determine the components’ masses. Astrometric measurements of both the ‘absolute’ positions in the sky and ‘relative’ positions of the two components to each other give direct access to the individual masses, which can be further

constrained with radial velocity monitoring. Such determinations have so far been made for a few dozen ultracool and brown dwarfs (e.g. Garcia et al. 2017; Dupuy & Liu 2017) and recently for the directly-imaged giant extrasolar planet β Pic b (Snellen & Brown 2018; Dupuy et al. 2019; Nielsen et al. 2020).

DENIS J063001.4–184014 (hereafter DE0630–18) was discovered as a late-type object by Phan-Bao et al. (2008) and classified as an M8.5 dwarf from its optical spectrum. It is part of a long-term astrometric monitoring campaign that makes use of the FORS2 optical camera mounted on the Very Large Telescope (Sahlmann et al. 2014). The discovery and characterisation of the binary orbit with an orbital period of 3.067 ± 0.006 years is described in Sahlmann et al. (2015a). Its proximity (~ 20 pc) and expected orbital separation should make it possible to resolve the components with infrared (IR) adaptive optics (AO) facilities on 10m-class telescopes. We

^{*} Based on observations made with ESO telescopes at the La Silla Paranal Observatory under programme IDs 086.C-0680, 088.C-0679, 090.C-0786, 092.C-0202.

[†] E-mail: jsahlmann@sciops.esa.int (JS)

therefore initiated corresponding follow-up observations in 2014.

Here, we present a detailed characterisation and dynamical mass determinations for the components in the DE0630–18 binary, made possible by a newly-obtained relative position measurement complemented by radial-velocity monitoring.

2 OBSERVATIONS AND DATA REDUCTION

2.1 Very Large Telescope/FORS2 astrometry

Here we do not present new FORS2 measurements, but we improve the accuracy of prior measurements by accounting for the updated camera calibrations presented in Lazorenko & Sahlmann (2017) and Lazorenko & Sahlmann (2019). These improvements include the use of Gaia catalogues (Gaia Collaboration et al. 2016) for external astrometric calibration and accounting for the subtle changes in the camera’s CCD chip locations.

For the conversion of our differential astrometric measurements to ICRS we used Gaia DR2 and applied the transformation procedure similar to that which we described in Lazorenko & Sahlmann (2018). As the comparison epoch, we adopted the average epoch $\bar{T} = 56249.955966$ MJD. Applying different polynomials for the transformation between FORS2 and DR2 and different criteria for elimination of outliers we derived solutions with different samples of common stars whose number varied between 21 to 57, just sufficient given the number of coefficients between 10 and 28 per axis. We derived an RMS of 4 mas for the ‘FORS2-DR2’ difference in positions, which exceeds its expected value of about 1 mas. Such a large RMS value is probably due to the different scale and relative rotation of the chips in the FORS2 detector, which was not introduced in the transformation model whereas the difference in zero points was taken into account as discussed in Lazorenko & Sahlmann (2019).

This transformation quality is however acceptable for the purpose of presenting the FORS2 positions in the ICRS after elimination of the geometric field deformation, its rotation, and the calibration of the pixel scale.

Table 1 contains the measured photocentre astrometry, which is given relative to the reference position $\alpha_0 = 97.507229262^\circ$, $\delta_0 = -18.672554677^\circ$ at time \bar{T} . We estimated that the actual accuracy of the reference position is ~ 5 mas, however the relative positions are about an order of magnitude more accurate.

2.2 Keck/NIRC2 LGS AO

We first observed DE0630–18 with the laser guide star adaptive optics (LGS AO) system at the Keck II telescope (Bouchez et al. 2004; Wizinowich et al. 2006; van Dam et al. 2006) on 2014 Mar 14 UT. We obtained data using the 9-hole non-redundant aperture mask installed in the filter wheel of NIRC2 (Tuthill et al. 2006) and analyzed these data using the same pipeline as in our previous work (e.g. Dupuy et al. 2009, 2015; Dupuy & Liu 2017). Interferograms taken in both the H and K bands, from the standard Mauna Kea Observatories (MKO) filters (Simons & Tokunaga 2002; Tokunaga et al. 2002), showed a significant detection of a binary. Fitting the closure phases derived from the higher-quality K -

band data gave a separation of 48 ± 3 mas, position angle (PA) of $302^\circ \pm 3^\circ$ and flux ratio of 1.88 ± 0.13 mag, where errors were computed using a Monte Carlo method that accounted for the measured closure phase errors. We observed DE0630–18 again on 2017 Mar 20 UT, this time in the MKO K_S band with a PSF calibrator observed immediately after the science target. This allowed us to measure more accurate binary parameters with a separation of 36.6 ± 1.3 mas, PA of $291.5 \pm 1.2^\circ$, and flux ratio of 1.74 ± 0.06 mag (Table 2).

The PA measured from the 2014 data (Table 3) is highly inconsistent ($\approx 10^\circ$ off, see Figure 5) compared to the prediction from the absolute astrometric orbit from FORS data. Such a discrepancy is unprecedented in our experience with other data obtained for similar binaries, even accounting for the lack of a PSF calibrator in 2014. Ultimately, we choose to exclude the 2014 measurements from our analysis, and the source of the PA discrepancy remains unknown.

2.3 High-resolution infrared spectroscopy

DE0630–18 was observed with the Keck II Near InfraRed Spectrometer (NIRSPEC; McLean et al. 2000) on four nights: 2016 November 16, 2017 February 6, 2017 March 22 and 2017 December 7 (UT). For each observation we used the N7 order-sorting filter and $0''.432$ -wide slit to obtain 2.00 – $2.39 \mu\text{m}$ spectra over orders 32–38 with $\lambda/\Delta\lambda = 20,000$ ($\Delta v = 15$ km/s) and dispersion of $0.315 \text{ \AA pixel}^{-1}$. Two dithered exposures of 1000 s (2016 Nov) or 600 s (2017 Feb, Mar and Dec) each were obtained, along with observations of the nearby A0 V star HD 49529 ($V = 8.09$). Flat field lamp, arc lamp and dark frame exposures were obtained at the start of each night for calibration. The NIRSPEC data were reduced and forward-modeled as described in Burgasser et al. (2015), Triaud et al. (2020) and Hsu et al. (in prep.), using a modified version of NIRSPEC Data Reduction Pipeline for reduction (Tran et al. 2016), telluric absorption models from Moehler et al. (2014), and the BT-Settl solar-metallicity atmosphere models (Allard et al. 2011) for the target spectrum. The analysis of the NIRSPEC data focused on the order 33, which covers both the CO $\nu = 2-0$ band at $2.29 \mu\text{m}$ and telluric CO and H_2O absorption used to refine the wavelength solution. Our forward-modeling method first fits a associated A0 V star spectrum to measure the instrumental line-spread function (LSF) modeled as a Gaussian broadening kernel, airmass, and precipitable water vapor (pwv). We then fit the stellar spectrum to a nine parameter model using a Markov Chain Monte Carlo (MCMC) algorithm, including four stellar parameters (effective temperature T_{eff} , surface gravity $\log g$, projected rotational velocity $v \sin i$, and radial velocity RV), two meteorological parameters (airmass and pwv), and three nuisance parameters (flux and wavelength offsets, and a noise scale factor). The rotational broadening profile assumes a limb-darkening coefficient of 0.6 (Gray 1992). We also fit out a 10th-order polynomial continuum correction at the end of each MCMC step. The MCMC was run with 50 walkers and 2,000 steps, with a sigma-clipping mask threshold of 2.5σ used to reject outlying pixels beyond the 1,000 step (less than 2% of the pixels were removed). Final fit parameters were determined from the mean and distribution of values in the final 800 steps, and a barycentric correction was applied to the inferred RV. We also re-evaluated the UVES data from Sahlmann et al. (2015b) following similar methods, focusing

Table 1. Individual FORS2 astrometric measurements of DE0630–18 for illustration. The complete table is available in electronic format. E is the epoch number and $\Delta\alpha_m^*$ and $\Delta\delta_m$ are the offsets in frame m relative to the reference position α_0, δ_0 at time \bar{T} . The four f coefficients are required to model DCR, as explained in Section 3.3.

E	m	t_m (MJD)	$\Delta\alpha_m^*$ (mas)	$\sigma_{\alpha_m^*}$ (mas)	$\Delta\delta_m$ (mas)	σ_{δ_m} (mas)	$f_{1,x,m}$	$f_{2,x,m}$	$f_{1,y,m}$	$f_{2,y,m}$
1	1	55537.27016	-496.105	0.738	841.776	0.825	0.08339	0.10735	0.10616	0.13666
1	2	55537.27077	-496.548	0.915	841.982	1.020	0.08701	0.11004	0.10635	0.13450
1	3	55537.27138	-496.512	0.865	842.772	0.960	0.09060	0.11259	0.10652	0.13237
1	4	55537.27200	-496.613	0.769	842.034	0.850	0.09418	0.11501	0.10667	0.13027
1	5	55537.27261	-495.803	0.749	842.397	0.847	0.09774	0.11732	0.10681	0.12820

Table 2. Results of Keck aperture masking (10 March 2017).

Sep.	(mas)	36.6 ± 1.3
PA	($^\circ$)	291.5 ± 1.2
ΔK_S	(mag)	1.74 ± 0.06

Table 3. Results of Keck aperture masking (13 March 2014). These data were not used in the analysis.

		H	K_s
Sep.	(mas)	48.1 ± 2.7	51.8 ± 2.2
PA	($^\circ$)	302.0 ± 2.4	304.3 ± 3.2
ΔMag	(mag)	N/A	1.88 ± 0.13
ΔMag	(mag)	1.87 ± 0.19	N/A

on the 819 nm Na I doublet, which falls in a relatively high S/N = 34 region of the observed data.

Table 4 summarizes the resulting atmospheric parameters and radial and rotational velocities determined by this analysis, while Figure 1 displays fits for one epoch each of the UVES and NIRSPEC data. Both the effective temperature ($T_{eff} = 2634 \pm 44$ K for UVES, $\langle T_{eff} \rangle = 2750 \pm 13$ K for NIRSPEC) and rotational velocity measurements ($v \sin i = 9.3 \pm 1.1$ km/s for UVES, $\langle v \sin i \rangle = 10.6 \pm 0.7$ km/s for NIRSPEC) of DE0630–18 are consistent across all epochs. The surface gravity shows a full dex discrepancy between the UVES ($\log g = 4.54 \pm 0.11$) and NIRSPEC ($\langle \log g \rangle = 5.49 \pm 0.01$) data, an issue previously noted in optical and infrared spectral modeling of the young eclipsing brown dwarf binary SPC J1510–2828AC (Triaud et al. 2020). The radial velocity measurements, which also vary significantly between epochs, is fortunately insensitive to the atmospheric parameters. We obtain statistically identical results if we constrain the fits to either the low or high values of $\log g$ inferred.

3 DATA ANALYSIS

3.1 Parallax correction

In the pre-Gaia work of Sahlmann et al. (2014) we determined the correction from relative to absolute parallax on the basis of a Galaxy model. Here, we can improve this by deriving the offset $\Delta\varpi$ for our relative parallaxes ϖ_F using the absolute Gaia DR2 parallaxes ϖ_G (Gaia Collaboration et al. 2018).

As described in Lazorenko et al. (2014), the astrometric reduction was made with different values of the model param-

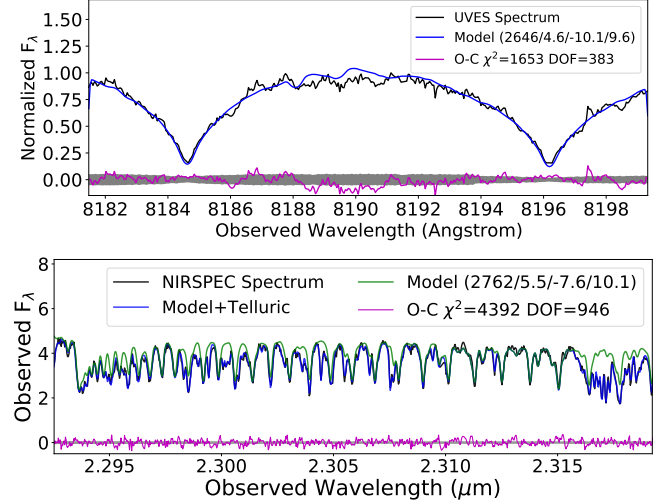


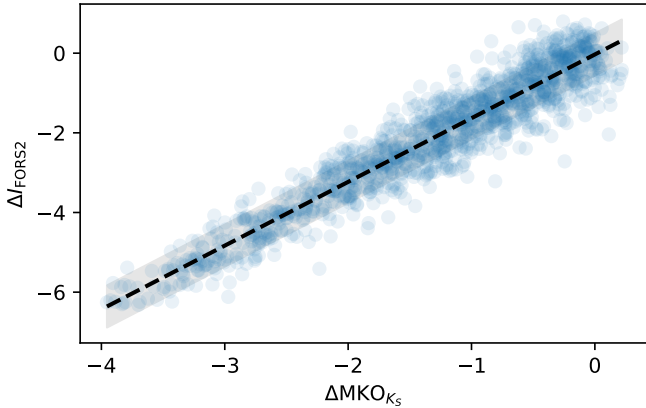
Figure 1. Forward model fits to the UVES spectrum (top) and NIRSPEC spectrum (bottom) of DE0630-18. Both panels show the data in black, the best-fit model in blue, and the difference spectrum in magenta (overlapping the 1σ uncertainty range in grey). The dashed lines in the UVES plot indicate masked-out telluric absorption features, while the green line in the NIRSPEC plot shows the model without telluric absorption included.

ter k , which determines the degree of the bivariate polynomial used to map individual frames to each other, and corresponding radii R of the reference field (up to $\sim 2'$ for $k = 16$). The final solution is taken as the weighted average of these individual solutions. Accordingly, we derived the parallax correction $\Delta\varpi$ using varying values of k and R by cross-matching the FORS2 stars and Gaia DR2 sources.

Many of the FORS2 reference stars lie at or beyond the faint end of Gaia DR2, thus were not used. For $k = 10$ and $k = 14$ we matched 32 and 63 reference stars, respectively, and derived $\Delta\varpi$ as the simple arithmetic average of differences $\varpi_G - \varpi_F$. After rejecting a few Gaia sources with abnormally large parallax uncertainties, we obtained corrections of $\Delta\varpi = 0.355 \pm 0.076$ mas and $\Delta\varpi = 0.331 \pm 0.059$ mas, respectively. These estimates are in good agreement and we adopted their average value of $\Delta\varpi = 0.343 \pm 0.068$ mas as the final estimate, which is also compatible with the initial estimate of 0.42 ± 0.04 mas in Sahlmann et al. (2014).

Table 4. Radial and Rotational Velocities from UVES & NIRSPEC Observations.

Instrument	MJD	Median S/N	T_{eff} (K)	$\log g$ (cm/s ²)	RV (km/s)	$v \sin i$ (km/s)
VLT/UVES	56568.36294	34	2634±44	4.54±0.11	-10.13±0.71	9.3±1.1
Keck/NIRSPEC	57708.49016	33	2735±7	5.49±0.02	-9.65±0.27	11.2±0.7
Keck/NIRSPEC	57790.27401	68	2758±7	5.49±0.01	-7.48±0.17	10.4±0.7
Keck/NIRSPEC	57834.22386	50	2750±5	5.49±0.01	-5.90±0.14	10.9±0.8
Keck/NIRSPEC	58094.49858	52	2752±5	5.49±0.01	-7.82±0.15	10.2±0.6

**Figure 2.** Synthesised magnitude differences in I_{FOR2} and MKO_{K_S} and the corresponding linear fit (dashed line).

3.2 Converting MKO K -band magnitude difference to the FORS2 I -band

To convert the photocentric FORS2 astrometry into barycentric astrometry a measurement of the magnitude difference in the FORS2 filter (ΔI_{FOR2}) is required. The Keck aperture mask observations yield a measurement of ΔK , which we needed to convert to ΔI_{FOR2} . We used the <https://github.com/BDNYC/BDNYCdb> database of ultracool spectra and the spectral energy distribution (SED) tool available at <https://github.com/hover2pi/sedkit> (Filippazzo et al. 2015) to generate synthetic magnitudes in the I_{FOR2} and K_S bands for a library of 141 M and L dwarfs, which includes 75 field objects and 66 low-gravity objects. For every source synthetic magnitudes were computed from its SED.

The 1230 pairwise combinations of sources with spectral types between M6 and L8 were used to estimate the magnitude differences shown in Figure 2. The relationship is well approximated by a straight line and we used the coefficients of the linear fit (not accounting for individual data point uncertainties) to convert the measured ΔK_S to ΔI_{FOR2} , and we added the residual RMS of the fit (0.53 mag) in quadrature to the uncertainty in ΔK_S . The measured value of $\Delta K_S = 1.74 \pm 0.06$ mag (Table 2) is then converted to $\Delta I_{\text{FOR2}} = 2.82 \pm 0.53$.

3.3 Combined astrometric model

The astrometric measurements of the target are $\alpha_m^* = \alpha_m \cos \delta$ and δ_m , corresponding to Right Ascension and Declination, respectively, in frame m at time t_m relative to the

reference frame of background stars. These are modeled with seven free parameters $\Delta\alpha_0^*$, $\Delta\delta_0$, μ_{α^*} , μ_{δ} , ϖ , d , and ρ as:

$$\begin{aligned} \alpha_m^* &= \Delta\alpha_0^* + \mu_{\alpha^*} t_m + \varpi \Pi_{\alpha,m} - \rho f_{1,x,m} - d f_{2,x,m} \\ \delta_m &= \Delta\delta_0 + \mu_{\delta} t_m + \varpi \Pi_{\delta,m} + \rho f_{1,y,m} + d f_{2,y,m}, \end{aligned} \quad (1)$$

where $\Delta\alpha_0^*$, $\Delta\delta_0$ are the coordinate offsets, μ_{α^*} , μ_{δ} are the proper motions, and the parallactic motion is expressed as the product of relative parallax ϖ and the parallax factors Π_{α} , Π_{δ} . The atmospheric differential chromatic refraction (DCR) is modelled with the free parameters ρ and d (Lazorenko et al. 2011; Sahlmann et al. 2014) and the coefficients f_1 and f_2 , where the latter are fully determined as a function of zenith angle, temperature, and pressure (Lazorenko 2006; Sahlmann et al. 2013, 2016). The DCR treatment does not involve the estimation of source colours, instead ρ and d are empirical free model parameters that correspond to the effective colour of the target relative to the average reference star.

The Keplerian orbit model adds an additional seven free parameters to the model. These are the eccentricity e , the argument of periastron ω , the orbital period P , the longitude of ascending node Ω , the orbital inclination i , the time of periastron passage T_P , and the semi-major axis of the photocentre orbit a_{phot} . Since we estimated the magnitude difference between the components in the filter bandpass in Section 3.2, we can relate the photocentre orbit size to the barycentre orbit size a_1 of the primary, similarly to Sahlmann et al. (2020). In the photocentre astrometry model that is applied to the FORS2 data we also include the astrometric nuisance parameters s_{α} and s_{δ} which can account for excess noise in the astrometry (Sahlmann et al. 2013).

The five available radial velocity measurements of the primary (Table 4) were modelled in a standard way. As discussed in Sahlmann et al. (2015b) we do not expect significant offsets between the UVES and NIRSPEC instruments and can therefore model those data jointly. Since the radial velocity variation of the primary is fully characterised by the orbital parameters above, the inclusion of radial velocities adds only the systemic velocity γ as free parameter.

Finally, we included the relative astrometric measurement from the resolved Keck aperture mask observation (Table 2). This allows us to determine model-independent masses because we can now directly adjust the primary and companion mass as free parameters (M_1 and M_2) instead of a_{phot} or a_1 . Apart from that, this step does not introduce any new parameters because the relative orbit is fully determined by the orbital parameters and the component masses.

3.4 MCMC analysis

We used the astrometric and orbital parameters from [Sahlmann et al. \(2015a\)](#) and reasonable guesses for the component masses and systemic radial velocity as starting values for a Markov Chain Monte Carlo (MCMC) analysis similar to that described in [Sahlmann et al. \(2020\)](#). We used the `emcee` package ([Foreman-Mackey et al. 2013](#)) to implement the MCMC and expressed the binary model using `pystrometry` ([Sahlmann 2019](#))¹ with the parameter vector θ composed of $\Delta\alpha_0^*$, $\Delta\delta_0$, ϖ , μ_{α^*} , μ_δ , ρ , d , P , e , ω , T_P , s_α , s_δ , M_1 , M_2 , i and Ω , γ .

The magnitude difference ΔI_{FOR2} and the parallax correction $\Delta\varpi$ are incorporated as Gaussian priors in the MCMC (see [Sahlmann et al. 2020](#)). Finally, the reference time T_{Ref} and the absolute coordinates enter the model as constants. The full model has 18 free parameters and two additional parameters constrained by priors.

3.5 Results

In our analyses we are always using all the individual frame data for the model fitting. For better visualisation of the results, however, we display only the epoch averages in the figures. Figure 3 shows the photocentric orbit of the binary in the sky and Figure 4 shows the same as a function of time. Figures 5 and 6 show the relative orbit with the Keck aperture mask measurement and Figure 7 shows the radial velocity curve from the UVES and NIRSPEC measurements.

Table 5 lists the adopted solution parameters determined as the median of the posterior distributions with 1 σ -equivalent confidence intervals. The results are unexpected in the sense that the primary mass of $0.052^{+0.009}_{-0.008} M_{\text{Sun}}$ is very low for an M8.5 dwarf and that the companion has essentially the same mass as its host. For comparison, the mean of seven dynamical masses for M8–M8.5 dwarfs measured by [Dupuy & Liu \(2017\)](#) is $0.090 M_{\text{Sun}}$, within 1σ of the *total* mass of the DE0630–18 system $0.104^{+0.013}_{-0.012} M_{\text{Sun}}$.

To estimate the dependency of this result on the empirically constrained Δmag parameter, we repeated the same analysis with larger magnitude differences, i.e. $\Delta I'_{\text{FOR2}} = \Delta I_{\text{FOR2}} + 1.5$ mag and $\Delta I''_{\text{FOR2}} = \Delta I_{\text{FOR2}} + 10$ mag. The results are reported in Tables 6 and 7, respectively, and we concluded that our main results do not depend on the exact value of the magnitude difference, because the derived masses agree within the uncertainties.

The overall fit quality (photocentre motion, relative separation, RV) is comparable and cannot be used to identify the more likely scenario among the following cases:

- Nominal ΔI_{FOR2} : This case corresponds to the ‘nominal’ optical flux ratio derived from the measured K -band flux ratio and the empirical relationship from Section 3.2. The primary mass and the companion mass are identical to each other within the errors. The observational data appear not to be fully compatible with the prior on ΔI_{FOR2} ($=2.8 \pm 0.5$ mag) since the corresponding posterior is shifted towards a higher value (the median deviates by +0.7 mag).

- $\Delta I'_{\text{FOR2}} = \Delta I_{\text{FOR2}} + 1.5$ mag: The primary is marginally more massive than the companion (Table 6). The

Table 5. Solution derived from the MCMC.

Parameter	Value
$\Delta\alpha_0^*$ (mas)	$574.59^{+0.17}_{-0.17}$
$\Delta\delta_0$ (mas)	$-902.60^{+0.08}_{-0.08}$
ϖ_{abs} (mas)	$51.34^{+0.09}_{-0.09}$
μ_{α^*} (mas yr ⁻¹)	$326.74^{+0.05}_{-0.05}$
μ_δ (mas yr ⁻¹)	$-502.65^{+0.03}_{-0.03}$
ρ (mas)	$19.91^{+0.71}_{-0.72}$
d (mas)	$-26.21^{+0.60}_{-0.60}$
P (day)	$1135.14^{+1.67}_{-1.74}$
P (yr)	$3.108^{+0.005}_{-0.005}$
Ω (°)	$111.65^{+0.12}_{-0.12}$
γ (m s ⁻¹)	$-10689.43^{+119.15}_{-123.47}$
λ_{ref} (°)	$-176.94^{+0.21}_{-0.20}$
$\sqrt{e} \sin \omega$ (°)	$-0.31^{+0.01}_{-0.01}$
$\sqrt{e} \cos \omega$ (°)	$0.29^{+0.01}_{-0.01}$
$\sqrt{M_2} \sin i$ (M_{Jup})	$7.35^{+0.32}_{-0.30}$
$\sqrt{M_2} \cos i$ (M_{Jup})	$-0.35^{+0.02}_{-0.02}$
s_α (mas)	$0.22^{+0.08}_{-0.10}$
s_δ (mas)	$0.16^{+0.10}_{-0.11}$
e (°)	$0.183^{+0.007}_{-0.008}$
ω (°)	$-46.76^{+1.80}_{-1.72}$
i (°)	$92.71^{+0.15}_{-0.15}$
T_P (day)	$56660.70^{+5.49}_{-5.17}$
a_{phot} (mas)	$23.55^{+0.06}_{-0.06}$
a_1 (mas)	$25.46^{+0.80}_{-0.61}$
a_{rel} (mas)	$51.42^{+2.05}_{-1.99}$
a_{rel} (AU)	$1.00^{+0.04}_{-0.04}$
M_2 (M_{Jup})	$54.11^{+4.81}_{-4.36}$
M_{tot} (M_{Sun})	$0.104^{+0.013}_{-0.012}$
M_1 (M_{Sun})	$0.052^{+0.009}_{-0.008}$
M_2 (M_{Sun})	$0.052^{+0.005}_{-0.004}$
Priors	
$\Delta\varpi$ (mas)	$0.34^{+0.07}_{-0.07}$
ΔI_{FOR2} (mag)	$3.54^{+0.43}_{-0.40}$

prior on $\Delta I'_{\text{FOR2}}$ ($=4.3 \pm 0.5$ mag) is better fulfilled and the posterior’s median deviates only by +0.3 mag.

- $\Delta I''_{\text{FOR2}} = \Delta I_{\text{FOR2}} + 10$ mag: In this case the companion is essentially dark and photocentre and barycentre motion coincide. Table 7 tabulates the median posterior values. The secondary is slightly less massive than in the previous case and the primary mass remains unchanged. The prior on $\Delta I''_{\text{FOR2}}$ ($=12.8 \pm 0.5$) is fulfilled.

4 DISCUSSION AND INTERPRETATION

4.1 Absolute Magnitudes and Colour

The 2MASS K_S -band magnitude of DE0630–18 is 11.46 ± 0.03 mag ([Cutri et al. 2003](#)). Our absolute paral-

¹ <https://github.com/Johannes-Sahlmann/pystrometry>

Table 6. Results with $\Delta I'_{\text{FORs2}} = \Delta I_{\text{FORs2}} + 1.5$ mag

Parameter	Value
γ (m s ⁻¹)	$-10552.14^{+95.51}_{-100.11}$
a_{phot} (mas)	$23.55^{+0.06}_{-0.06}$
a_1 (mas)	$24.31^{+0.41}_{-0.29}$
a_{rel} (mas)	$51.36^{+1.96}_{-1.90}$
a_{rel} (AU)	$1.00^{+0.04}_{-0.04}$
M_{tot} (M_{Sun})	$0.104^{+0.012}_{-0.011}$
M_1 (M_{Sun})	$0.055^{+0.008}_{-0.007}$
M_2 (M_{Sun})	$0.049^{+0.004}_{-0.004}$
$\Delta I'_{\text{FORs2}}$ (mag)	$4.56^{+0.50}_{-0.47}$

Table 7. Results with $\Delta I''_{\text{FORs2}} = \Delta I_{\text{FORs2}} + 10$ mag

Parameter	Value
γ (m s ⁻¹)	$-10447.03^{+87.16}_{-90.73}$
a_{phot} (mas)	$23.54^{+0.06}_{-0.06}$
a_1 (mas)	$23.54^{+0.06}_{-0.06}$
a_{rel} (mas)	$51.17^{+1.85}_{-1.88}$
a_{rel} (AU)	$1.00^{+0.04}_{-0.04}$
M_{tot} (M_{Sun})	$0.103^{+0.011}_{-0.011}$
M_1 (M_{Sun})	$0.055^{+0.008}_{-0.007}$
M_2 (M_{Sun})	$0.047^{+0.003}_{-0.003}$
$\Delta I''_{\text{FORs2}}$ (mag)	$12.81^{+0.53}_{-0.52}$

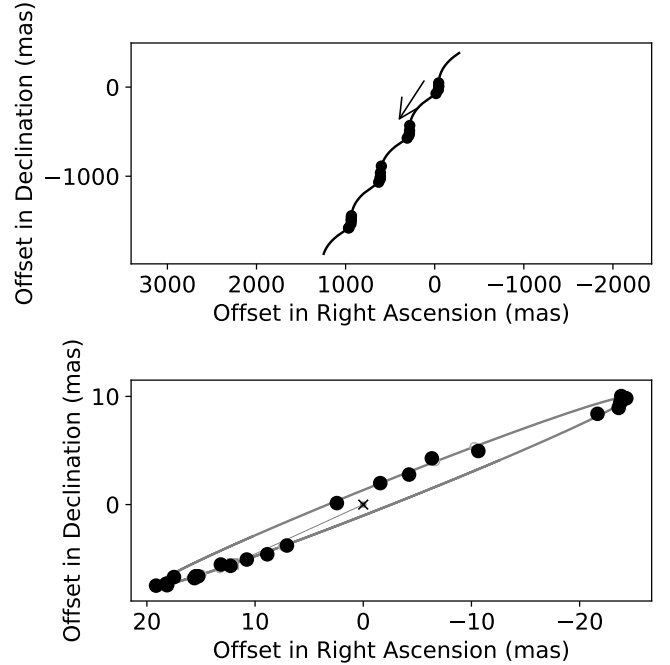
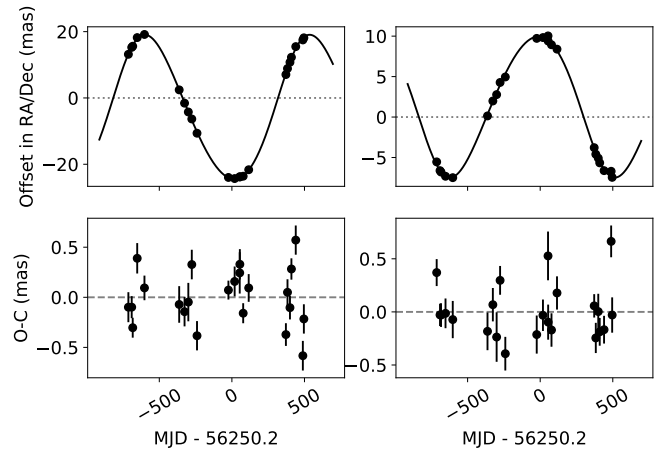
lax implies a distance of 19.478 ± 0.034 pc, for a system absolute magnitude of $M_{K_s} = 10.01 \pm 0.03$ mag. Applying our measured relative magnitude of 1.74 ± 0.06 mag in the Keck/NIRC2 K_s -band filter, we infer component absolute magnitudes $M_{K_s} = 10.17 \pm 0.03$ mag for the primary and $M_{K_s} = 12.04 \pm 0.11$ mag for the secondary.

The primary is thus pretty much spot on for an M8.5. The secondary is 1.7 mag fainter in K , thus faint enough that it could correspond to an L5-L6 dwarf.

While the absolute brightness of DE0630–18 is consistent with expectations, the near-infrared colour and spectrum of this source are unusual. Its 2MASS $J - K_s = 1.22 \pm 0.16$ mag is relatively red compared to optically-classified M8 dwarfs ($\langle J - K_s \rangle = 1.03 \pm 0.17$ mag; Schmidt et al. 2015). This is clearly apparent in low-resolution near-infrared spectral data of DE0630–18 from Bardalez Gagliuffi et al. (2019), which has a consistent spectrophotometric color of $J - K_s = 1.26$ and diverges from the M8 dwarf standard VB 10 beyond $1.3 \mu\text{m}$ (Figure 8). We found even worse agreement with intermediate-gravity (INT-G or β) and very low-gravity (VL-G or γ) M8 standards, consistent with the field gravity classification based on the methods of Allers & Liu (2013).

4.2 Spectral binary fit

The most obvious explanation for this deviation is the contribution of secondary light to the blended-light spectrum. We

**Figure 3.** Best fit model of FORS2 parallax/proper motion (top) and photocentric orbit (bottom). The black circles indicate the epoch-averaged FORS2 measurements with uncertainties that are smaller than the symbol size.**Figure 4.** Top: Best fit model of FORS2 photocentric orbit as a function of time. Bottom: O-C residuals. The fit quality (0.27 mas RMS) is marginally deteriorated compared to the astrometry-only fit of Sahlmann et al. (2015a, Fig. 4, 0.24 mas RMS).

compared the spectrum of DE0630–18 to three sets of binary templates constructed from data drawn from the SpeX Prism Library (Burgasser 2014). We created two “field” template sets, using high signal-to-noise spectra ($S/N \geq 100$) of 285 M7-M9 dwarfs for the primaries, and either $S/N \geq 75$ spectra of 215 L1-L5 dwarfs (“early-type secondary”) or $S/N \geq 30$ spectra of 79 L5-L9 dwarfs (“late-type secondary”) for the secondaries. We created a third “young” binary template set by combining $S/N \geq 70$ spectra of 104 M7-M9 and

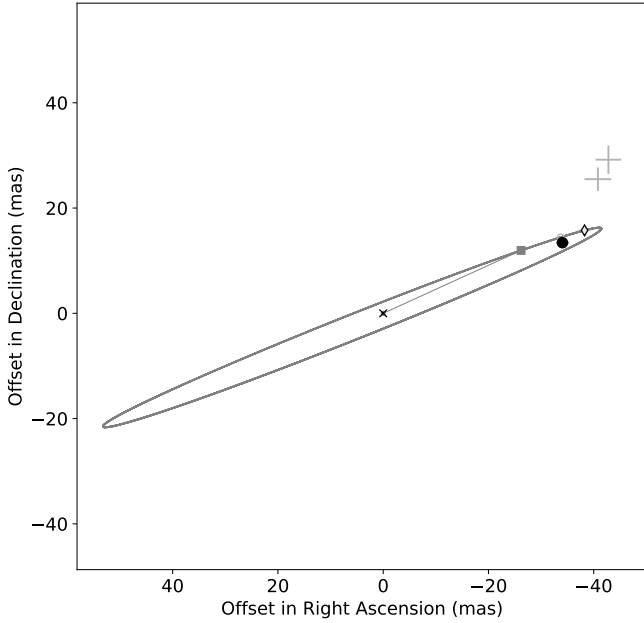


Figure 5. Best fit model of the relative orbit. The black symbol shows the single Keck aperture masking measurement. The grey crosses mark the discarded 2014 measurements and the diamond marks their expected position.

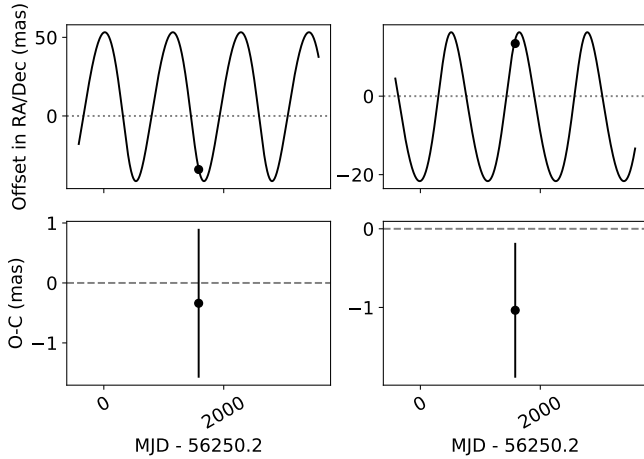


Figure 6. Top: Best fit model of relative orbit as a function of time. Bottom: O-C residuals. The black symbol shows the single used Keck aperture masking measurement.

$S/N \geq 30$ spectra of 78 L3-L7 intermediate-gravity dwarfs. The gravity classifications of all templates were confirmed using the index-based scheme defined in Allers & Liu (2013). The spectral components of all template binaries were scaled to match the observed $\Delta K_s = 1.74$ mag from the aperture masking observations. The bottom three panels in Figure 8 show that field-gravity and intermediate-gravity binary templates reproduce the spectrum of DE0630–18 significantly better than the single standards. The field template binaries provide better fits than the intermediate-gravity template, which shows low-level deviations across the $0.9\text{--}1.3\ \mu\text{m}$ region, likely arising from mismatches in FeH band and atomic

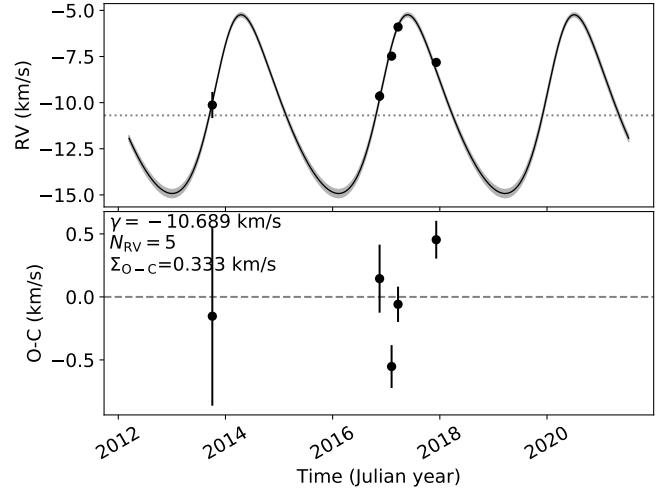


Figure 7. Top panel: RV measurements (black symbols; the UVES data are from ~ 2014 and the remaining data are from NIRSPEC) and the best-fit orbit (solid curve). The grey area corresponds to the 1-sigma equivalent range of RV values filled by random draws from the posteriors. The systemic velocity is shown by the horizontal dashed line. Bottom panel: O-C residuals.

line strengths. The field templates with early-type secondaries ($\sim L2$) and late-type secondaries ($\sim L7$) provide similar agreement, so these fits are insufficient to firmly determine the classification of the secondary, although the latter case would be consistent with an unresolved secondary component (see Section 4.5).

4.3 Moving group membership

Using our proper motion, parallax, and RV determinations of DE0630–18 we probed kinematic membership in known moving groups using the BANYAN tool² (Gagné et al. 2018), which yielded a 99.9 % probability that DE0630–18 is a field object, thus there is no indication that it belongs to any of the considered young moving groups.

4.4 Lithium

Despite the low masses in the system, we do not find signs of lithium absorption in our UVES spectrum (Sahlmann et al. 2015a). We could place an upper limit of $\sim 1\text{\AA}$ on the Li I equivalent width (pEW) by visual inspection of simulations that consisted of injecting a spectral signature in the observed spectrum.

At ages younger than 200 Myr, the lithium depletion boundary is located at $M_{K_s} < 10.0$ mag (Dahm 2015) and consequently the lack of lithium detection in the primary indicates an age older than 200 Myr for the system. Should the age of the system be older than 200 Myr but younger than 700 Myr, the secondary could have preserved its initial lithium content. Considering for example the case that the secondary would have a similar Li I equivalent width as that of Hya12, an L4 member in the Hyades cluster with a pEW

² <http://www.exoplanetes.umontreal.ca/banyan>

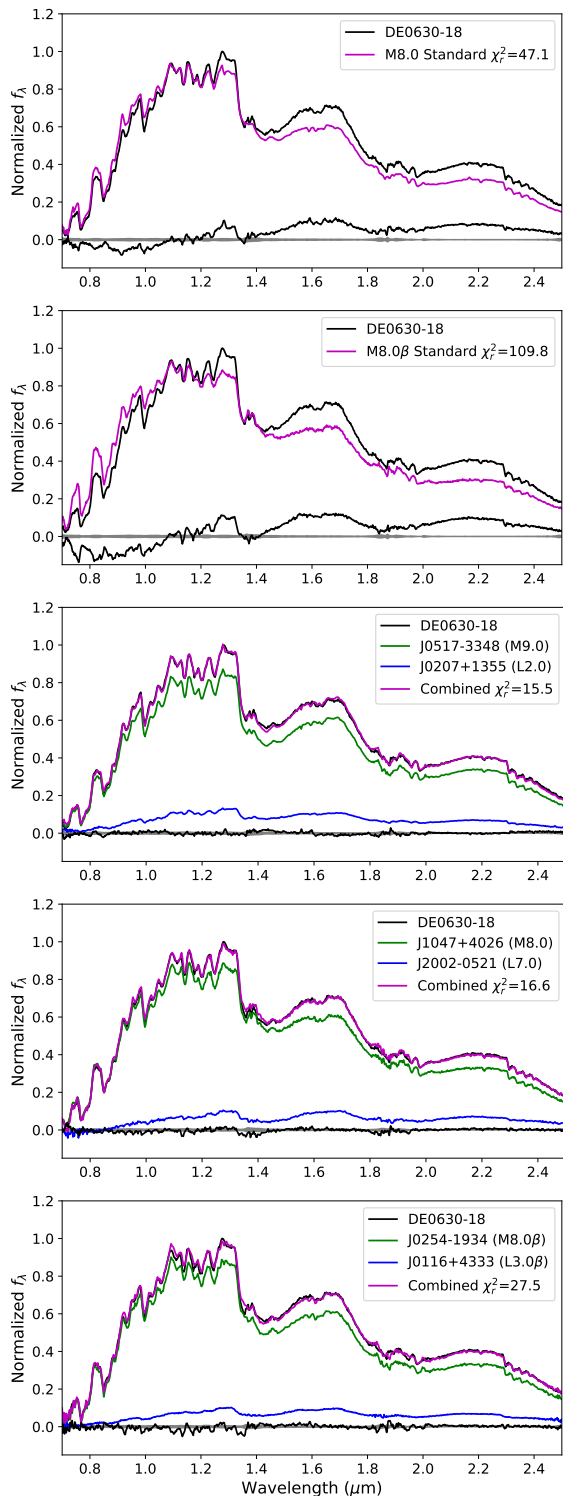


Figure 8. From top to bottom, comparison of DE0630–18’s near-infrared spectrum (black lines) to various spectral templates (magenta lines): the M8 dwarf standard VB 10 (data from [Bardalez Gagliuffi et al. 2014](#)), the intermediate-gravity M8 β standard 2MASS J0019262+461407 (data from [Bardalez Gagliuffi et al. 2019](#)), the best-fit blended-light spectrum composed of a field gravity late-M primary (green line) and early-L secondary (blue line), the best-fit blended-light spectrum composed of a field gravity late-M primary (green line) and late-L secondary (blue line), and the best-fit blended-light spectrum composed of an intermediate-gravity late-M primary (green line) and L dwarf secondary (blue line). The blended-light templates are all scaled to a relative magnitude of $\Delta K_s = 1.74$. All panels show the $\pm 3\sigma$ uncertainties of the DE0630–18 spectrum (grey region) and difference spectrum (black line) at a median value of zero.

of 8.5 Å and an age of 650 Myr ([Martín et al. 2018](#)). If the primary contributes only flux, and no lithium absorption at all, the lithium feature from the secondary would be diluted by a factor of 15.8 considering a magnitude difference of 3 at 670.8 nm. The lithium feature in the secondary could then appear as a moderately weak absorption with pEW of about 0.54 Å, which could not be detected in our UVES spectrum.

We note that a diluted lithium feature from the secondary could be clearly detected in a reasonable exposure time of about 1 hour with a similar instrumental setup as that used for observing Hya12 with OSIRIS at the 10.4-meter Gran Telescopio de Canarias. It would be interesting to attempt such observation in the future in order to derive additional constraints on the age of the system. DE0630–18 may thereby become a member of the exclusive club of brown dwarf binaries for which dynamical masses and lithium depletion factors can be determined for each component, joining systems such as GJ 569B ([Zapatero Osorio et al. 2005](#)).

4.5 Interpretation

We have shown that no matter what the true ΔI is, the primary mass remains very low. Assuming the primary is single, its $M_{K_s} \sim 10.0$ mag implies $\log L_{\text{bol}} \sim -3.3$ dex, and a mass of $0.055 M_{\text{sun}}$ implies an age of ~ 150 Myr from the SM08 models. This is quite young for being FLD-G, but not unprecedented ([Aller et al. 2016](#)). The problem is the companion. If it is single, its inferred $M_{K_s} \sim 12$ mag implies $\log L_{\text{bol}} \sim -4.1$ dex, and thus would have a mass of $\sim 0.030 M_{\text{sun}}$ according to the SM08 models. This is inconsistent with the measured component mass, as well as the combined mass of the system, $\sim 0.104 M_{\text{sun}}$. Moreover, an equal mass system is in tension with the distinct spectral morphologies inferred from the blended light infrared spectrum.

The empirical mass- M_{K_s} relationship of [Mann et al. \(2019\)](#) for $M_{K_s} \sim 10.0$ mag predicts a primary mass of $0.086 M_{\text{sun}}$ ($\pm 3\%$), which is also well above the astrometric mass measured here. Even if this higher mass is correct, the total mass of the system we require a secondary mass of only $0.030 M_{\text{sun}}$, again below its measured astrometric mass and consistent with a ~ 150 Myr-old system based on evolutionary models.

One scenario that could resolve these discrepancies is for DE0630–18 to be a young triple system with an unresolved binary secondary. For an age of ~ 150 Myr, a primary of mass $0.055 M_{\text{sun}}$ and an unresolved secondary with components each of mass $\sim 0.025 M_{\text{sun}}$ would be consistent with all observations. These include the measured masses, the relative K -band magnitudes and component types (primary $T_{\text{eff}} \approx 2400$ K, secondary $T_{\text{eff}} \approx 1400$ K), the absence of strong intermediate-gravity features in infrared spectra, and the lack of Li I in the combined-light optical spectra. Substellar triples are already been identified in both the field and young moving groups ([Bouy et al. 2005](#); [Radigan et al. 2013](#); [Dupuy & Liu 2017](#); [Triaud et al. 2020](#)), and DE0630–18 may represent such a system at an early stage of evolution.

5 CONCLUSIONS

We presented individual mass determinations of the components in the DE0630–18 system on the basis of photocentre

astrometry monitoring, one relative position determination, and a five radial velocity observations. The results indicate two components with nearly identical masses firmly in the substellar regime. This result does not depend on the precise value of the components' magnitude difference in the optical, which is the only quantity that we did not measure directly but that we had to determine empirically.

Given the object's M8.5 optical spectral type and absolute magnitude, but lack of Li I absorption, this suggests that DE0630–18 is a relatively young system (~ 200 Myr). The measured relative magnitude and distinct spectral components inferred from analysis of its combined-light spectrum are at tension with the mass measurements, but can be resolved if the secondary of DE0630–18 is itself an unresolved binary.

Additional observations, e.g. a second relative separation measurement and spectroscopy targeting the diluted lithium signature from the secondary, would be beneficial in clarifying the properties of DE0630–18.

ACKNOWLEDGEMENTS

This research made use of the databases at the Centre de Données astronomiques de Strasbourg (<http://cds.u-strasbg.fr>); NASA's Astrophysics Data System Service (http://adsabs.harvard.edu/abstract_service.html), the paper repositories at arXiv; the SpeX Prism Libraries (<http://www.browndwarfs.org/spexprism>); and ASTROPY, a community-developed core Python package for Astronomy (Astropy Collaboration et al. 2013). The authors also made use of SCIPY (Jones et al. 2001), NUMPY (Oliphant 2007), IPYTHON (Pérez & Granger 2007), and MATPLOTLIB (Hunter 2007). This work has made use of data from the ESA space mission *Gaia* (<http://www.cosmos.esa.int/gaia>), processed by the *Gaia* Data Processing and Analysis Consortium (DPAC, <http://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for the DPAC has been provided by national institutions, in particular the institutions participating in the *Gaia* Multilateral Agreement. This research was funded by the Ministerio de Economía y Competitividad and the Fondo Europeo de Desarrollo Regional (FEDER) under grant AYA2015-69350-C3-1-P. AJB acknowledges funding support from the National Science Foundation under award No. AST-1517177. The material is based upon work supported by the National Aeronautics and Space Administration under Grant No. NNX15AI75G. M. C. L. acknowledges support from National Science Foundation grant AST-1518339. The authors recognize and acknowledge the very significant cultural role and reverence that the summit of Mauna Kea has always had within the indigenous Hawaiian community. We are most fortunate and grateful to have the opportunity to conduct observations from this mountain.

DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding author.

REFERENCES

- Allard F., Homeier D., Freytag B., 2011, in Johns-Krull C., Browning M. K., West A. A., eds, *Astronomical Society of the Pacific Conference Series Vol. 448, 16th Cambridge Workshop on Cool Stars, Stellar Systems, and the Sun*. p. 91
- Aller K. M., et al., 2016, *ApJ*, 821, 120
- Allers K. N., Liu M. C., 2013, *ApJ*, 772, 79
- Astropy Collaboration et al., 2013, *A&A*, 558, A33
- Bardalez Gagliuffi D. C., et al., 2014, *ApJ*, 794
- Bardalez Gagliuffi D. C., et al., 2019, *ApJ*, 883, 205
- Bouchez A. H., et al., 2004, in Bonaccini Calia D., Ellerbroek B. L., Ragazzoni R., eds, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series Vol. 5490, Advancements in Adaptive Optics*. pp 321–330
- Bouy H., Martín E. L., Brandner W., Bouvier J., 2005, *AJ*, 129, 511
- Burgasser A. J., 2014, in *Astronomical Society of India Conference Series*. pp 7–16 ([arXiv:1406.4887](https://arxiv.org/abs/1406.4887))
- Burgasser A. J., et al., 2015, *AJ*, 149, 104
- Cutri R. M., et al., 2003, *2MASS All Sky Catalog of point sources.. NASA/IPAC*
- Dahm S. E., 2015, *ApJ*, 813, 108
- Dupuy T. J., Liu M. C., 2017, *ApJS*, 231, 15
- Dupuy T. J., Liu M. C., Ireland M. J., 2009, *ApJ*, 699, 168
- Dupuy T. J., Liu M. C., Leggett S. K., Ireland M. J., Chiu K., Golimowski D. A., 2015, *ApJ*, 805, 56
- Dupuy T. J., Brandt T. D., Kratter K. M., Bowler B. P., 2019, *ApJ*, 871, L4
- Filippazzo J. C., Rice E. L., Faherty J., Cruz K. L., Van Gordon M. M.,Looper D. L., 2015, *ApJ*, 810, 158
- Foreman-Mackey D., Hogg D. W., Lang D., Goodman J., 2013, *PASP*, 125, 306
- Gagné J., et al., 2018, *ApJ*, 856, 23
- Gaia Collaboration et al., 2016, *A&A*, 595, A2
- Gaia Collaboration Brown A. G. A., Vallenari A., Prusti T., de Bruijne J. H. J., Babusiaux C., Bailer-Jones C. A. L., 2018, preprint
- Garcia E. V., et al., 2017, *ApJ*, 846, 97
- Gray D. F., 1992, *The observation and analysis of stellar photospheres.. Vol. 20*
- Hunter J. D., 2007, *Computing in Science and Engineering*, 9
- Jones E., Oliphant T., Peterson P., et al., 2001, *SciPy: Open source scientific tools for Python*
- Lazorenko P. F., 2006, *A&A*, 449, 1271
- Lazorenko P. F., Sahlmann J., 2017, *A&A*, 606, A90
- Lazorenko P. F., Sahlmann J., 2018, *A&A*, 618, A111
- Lazorenko P. F., Sahlmann J., 2019, *A&A*, 629, A113
- Lazorenko P. F., et al., 2011, *A&A*, 527, A25+
- Lazorenko P. F., Sahlmann J., Ségransan D., Martín E. L., Mayor M., Queloz D., Udry S., 2014, *A&A*, 565, A21
- Liu M. C., Dupuy T. J., Ireland M. J., 2008, *ApJ*, 689, 436
- Mann A. W., et al., 2019, *ApJ*, 871, 63
- Martín E. L., Lodieu N., Pavlenko Y., Béjar V. J. S., 2018, *ApJ*, 856, 40
- McLean I. S., Graham J. R., Becklin E. E., Figer D. F., Larkin J. E., Levenson N. A., Teplitz H. I., 2000, in *SPIE*. pp 1048–1055
- Moehler S., et al., 2014, *A&A*, 568, A9
- Nielsen E. L., et al., 2020, *AJ*, 159, 71
- Oliphant T. E., 2007, *Computing in Science and Engineering*, 9
- Pérez F., Granger B. E., 2007, *Computing in Science and Engineering*, 9
- Phan-Bao N., et al., 2008, *MNRAS*, 383, 831
- Radigan J., Jayawardhana R., Lafrenière D., Dupuy T. J., Liu M. C., Scholz A., 2013, *ApJ*, 778, 36
- Sahlmann J., 2019, *Johannes-Sahlmann/pystrometry*, [doi:10.5281/zenodo.3515526](https://doi.org/10.5281/zenodo.3515526), <https://doi.org/10.5281/zenodo.3515526>

[zenodo.3515526](https://zenodo.org/record/3515526)

- Sahlmann J., Lazorenko P. F., Ségransan D., Martín E. L., Queloz D., Mayor M., Udry S., 2013, *A&A*, 556, A133
- Sahlmann J., Lazorenko P. F., Ségransan D., Martín E. L., Mayor M., Queloz D., Udry S., 2014, *A&A*, 565, A20
- Sahlmann J., Lazorenko P. F., Ségransan D., Martín E. L., Mayor M., Queloz D., Udry S., 2015a, *A&A*, 577, A15
- Sahlmann J., et al., 2015b, *A&A*, 579, A61
- Sahlmann J., Lazorenko P. F., Bouy H., Martín E. L., Queloz D., Ségransan D., Zapatero Osorio M. R., 2016, *MNRAS*, 455, 357
- Sahlmann J., et al., 2020, *MNRAS*, 495, 1136
- Schmidt S. J., Hawley S. L., West A. A., Bochanski J. J., Davenport J. R. A., Ge J., Schneider D. P., 2015, *AJ*, 149, 158
- Simons D. A., Tokunaga A., 2002, *PASP*, 114, 169
- Snellen I. A. G., Brown A. G. A., 2018, *Nature Astronomy*
- Tokunaga A. T., Simons D. A., Vacca W. D., 2002, *PASP*, 114, 180
- Tran H. D., et al., 2016, in *Observatory Operations: Strategies, Processes, and Systems VI*. p. 99102E
- Triaud A. H. M. J., et al., 2020, *Nature Astronomy*,
- Tuthill P., et al., 2006, in *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*. p. 62723A
- Wizinowich P. L., et al., 2006, *PASP*, 118, 297
- Zapatero Osorio M. R., Lane B. F., Pavlenko Y., Martín E. L., Britton M., Kulkarni S. R., 2004, *ApJ*, 615, 958
- Zapatero Osorio M. R., Martín E. L., Lane B. F., Pavlenko Y., Bouy H., Baraffe I., Basri G., 2005, *Astronomische Nachrichten*, 326, 948
- van Dam M. A., et al., 2006, *PASP*, 118, 310

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.